

Testing a Fast Dynamical Indicator: The MEGNO

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Abstract

To investigate non-linear dynamical systems, like for instance artificial satellites, Solar System, exoplanets or galactic models, it is necessary to have at hand several tools, such as a reliable dynamical indicator.

The aim of the present work is to test a relatively new fast indicator, the Mean Exponential Growth factor of Nearby Orbits (MEGNO), since it is becoming a widespread technique for the study of Hamiltonian systems, particularly in the field of dynamical astronomy and astrodynamics, as well as molecular dynamics.

In order to perform this test we make a detailed numerical and statistical study of a sample of orbits in a triaxial galactic system, whose dynamics was investigated by means of the computation of the Finite Time Lyapunov Characteristic Numbers (FT-LCNs) by other authors.

Keywords:

Non-Linear Dynamics, Chaos, Lyapunov Characteristic Number, MEGNO

1. Introduction

In the present work we accomplish an exhaustive study of the MEGNO when applied to a given sample of orbits in a triaxial galactic potential studied by [1]. In that work, the authors use a well-known tool, the Lyapunov

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Characteristic Numbers (see e.g. [2]), to identify the character of the selected orbits in order to classify them as regular or chaotic.

The MEGNO is introduced by [3] and, in [4], this technique is formalized and its application extended to discrete Hamiltonian systems like maps; also a generalization of the MEGNO is introduced therein. This tool has become of widespread use for studying several astronomical problems as well as many other Hamiltonian systems (see, e.g., [5, 6, 7, 8, 9, 10, 11, 12, 13, 14, 15, 16, 17, 18]).

In [4] and [19], the MEGNO succeed in furnishing a clear insight of the global structure of the phase space of simple multidimensional Hamiltonian systems, providing a clear picture of the resonant network as well as the regular and chaotic domains.

Herein instead, a far more complex non-linear system is addressed that reproduces many characteristics of real elliptical galaxies, namely, the one introduced by [1]. This model will be used as the scenario for a detailed comparison between the MEGNO and the Lyapunov Characteristic Numbers and even the Fast Lyapunov Indicator (FLI) introduced by [20].

There are many efficient dynamical indicators, some of them based on deviation vector(s), for instance, the spectra of stretching numbers, helicity and twist angles, the computation of the alignment indices introduced by Skokos, the Relative finite time Lyapunov Indicator (RLI) and also the Average Power Law Exponent (APLE), a technique recently developed [21, 22, 23, 24], and others based on spectral analysis, such as [25], the Frequency Map Analysis [26, 27], the one due to Sidlichovský and Nesvorný [28], and perhaps the latest one, the FMI (Frequency modulation indicator) [29]. However the present paper is devoted to accomplish a thorough test of the MEGNO, pointing out not only its advantages but its drawbacks as well. Therefore herein we just focus our attention on an exhaustive comparison of the MEGNO against the Lyapunov Characteristic Numbers (and eventually the FLI), since without any doubt, the latter is the most widespread tool in, at least, the last forty years, and it is still being used by many authors. On the other hand, as far as we know, a full test of the MEGNO when applied to a non-linear somewhat realistic Hamiltonian system has not been performed yet.

2. The Mean Exponential Growth factor of Nearby Orbits (MEGNO)

In this section we summarize the main features of the MEGNO (described in detail in [4]). This is an alternative tool to explore the phase space which

belongs to the class of the so-called fast indicators.

Let $H(\mathbf{p}, \mathbf{q})$ with $\mathbf{p}, \mathbf{q} \in \mathbb{R}^N$ be an N -dimensional Hamiltonian, that we suppose autonomous just for the sake of simplicity. Introducing the following notation:

$$\mathbf{x} = (\mathbf{p}, \mathbf{q}) \in \mathbb{R}^{2N}, \quad \mathbf{v} = (-\partial H/\partial \mathbf{q}, \partial H/\partial \mathbf{p}) \in \mathbb{R}^{2N},$$

the equations of motion can be written in a simple way like

$$\dot{\mathbf{x}} = \mathbf{v}(\mathbf{x}). \quad (1)$$

Let $\gamma(\mathbf{x}_0; t)$ be an arc of an orbit of the flow (1) over a compact energy surface: $M_h \subset \mathbb{R}^{2N}$, $M_h = \{\mathbf{x} : H(\mathbf{p}, \mathbf{q}) = h\}$ with $h = \text{constant}$, then

$$\gamma(\mathbf{x}_0; t) = \{\mathbf{x}(t'; \mathbf{x}_0) : \mathbf{x}_0 \in M_h, 0 \leq t' < t\}.$$

We can gain fundamental information about the Hamiltonian flow in the neighborhood of any orbit γ through the largest Lyapunov Characteristic Number (LCN) defined as:

$$\sigma(\gamma) = \lim_{t \rightarrow \infty} \sigma_1(\gamma(\mathbf{x}_0; t)), \quad \sigma_1(\gamma(\mathbf{x}_0; t)) = \frac{1}{t} \ln \left[\|\vec{\delta}\gamma(\mathbf{x}_0; t)\| \right], \quad (2)$$

with $\vec{\delta}\gamma(\mathbf{x}_0; t)$ an ‘‘infinitesimal displacement’’ from γ at time t , where $\|\cdot\|$ is some norm. The fact that the LCN measures the mean exponential rate of divergence of nearby orbits it is clearly understood when Eq. (2) is written in an integral fashion:

$$\sigma(\gamma) = \lim_{t \rightarrow \infty} \frac{1}{t} \int_0^t \frac{\dot{\delta}\gamma(\mathbf{x}_0; t')}{\delta\gamma(\mathbf{x}_0; t')} dt' = \overline{\left(\dot{\delta}/\delta \right)}, \quad (3)$$

where $\delta \equiv \|\vec{\delta}\|$, $\dot{\delta} \equiv d\delta/dt = \vec{\delta} \cdot \vec{\delta}/\|\vec{\delta}\|$, and the bar denotes time average. Also, the tangent vector $\vec{\delta}$ satisfies the variational equation

$$\dot{\vec{\delta}} = \Lambda(\gamma(\mathbf{x}_0; t)) \cdot \vec{\delta},$$

where Λ is the Jacobian matrix associated with the vector field \mathbf{v} .

Now we are in a position to introduce the MEGNO, $Y(\gamma(\mathbf{x}_0; t))$, through the expression:

$$Y(\gamma(\mathbf{x}_0; t)) = \frac{2}{t} \int_0^t \frac{\dot{\delta}\gamma(\mathbf{x}_0; t')}{\delta\gamma(\mathbf{x}_0; t')} t' dt',$$

which is related with the integral in Eq. (3); i.e., in case of an exponential increase of δ , $\delta\gamma(\mathbf{x}_0; t) = \delta\gamma(\mathbf{x}_0; t_0) \cdot \exp(\lambda t)$, the quantity $Y(\gamma(\mathbf{x}_0; t))$ can be considered as a weighted variant of the integral in Eq. (3). Instead of using the instantaneous rate of increase, λ , we average the logarithm of the growth factor, $\ln[\delta\gamma(\mathbf{x}_0; t)/\delta\gamma(\mathbf{x}_0; t_0)] = \lambda t$.

Let us describe the MEGNO's asymptotic behavior to exhibit its ability to give a clean idea of the character of orbits. Firstly, consider the case of orbits on irrational tori for a non-isochronous system. As it is shown in [4], for quasi-periodic orbits, γ_q , the temporal evolution of $Y(\gamma_q(\mathbf{x}_0; t))$ is given by

$$Y(\gamma_q(\mathbf{x}_0; t)) \approx 2 - \frac{\ln(1 + \lambda_q t)^2}{\lambda_q t} + O(\gamma_q(\mathbf{x}_0; t)),$$

where λ_q is the linear rate of divergence around γ_q and O is a null average oscillating term. Accordingly to this formula, the $\lim_{t \rightarrow \infty} Y(\gamma_q(\mathbf{x}_0; t))$ does not exist, but on introducing a time average

$$\bar{Y}(\gamma_q(\mathbf{x}_0; t)) \equiv \frac{1}{t} \int_0^t Y(\gamma_q(\mathbf{x}_0; t')) dt',$$

it can be found that

$$\bar{Y}(\gamma_q) \equiv \lim_{t \rightarrow \infty} \bar{Y}(\gamma_q(\mathbf{x}_0; t)) = 2.$$

Then, for quasi-periodic motion, $\bar{Y}(\gamma)$ is a fixed constant, independent of γ .

When taking irregular orbits γ_i , i.e. orbits on some stochastic layer, for which $\delta\gamma_i(\mathbf{x}_0; t) \approx \delta\gamma(\mathbf{x}_0; t_0) \cdot \exp(\sigma_i t)$, σ_i being the LCN of γ_i , the temporal evolution of the MEGNO will be given by:

$$Y(\gamma_i(\mathbf{x}_0; t)) \approx \sigma_i t + \tilde{O}(\gamma_i(\mathbf{x}_0; t)),$$

with \tilde{O} some bounded amplitude and null average oscillating term (see [4]). On averaging over a sufficiently large interval we have:

$$\bar{Y}(\gamma_i(\mathbf{x}_0; t)) \approx \frac{\sigma_i}{2} t, \quad t \rightarrow \infty.$$

Therefore, in the case of chaotic orbits, not only $Y(\gamma_i(\mathbf{x}_0; t))$ but also $\bar{Y}(\gamma_i(\mathbf{x}_0; t))$ grow linearly with time, with a slope equal to the LCN of the orbit or one half of it, respectively. Wherever the phase space has a hyperbolic structure,

\bar{Y} will indefinitely grow with time. Otherwise, it will approach a constant value, even in the degenerated case in which δ grows with some power of t , e.g. n , for which $\bar{Y} \rightarrow 2n$ when $t \rightarrow \infty$.

We notice that the temporal evolution of the MEGNO can be briefly described in a suitable and unique expression for all kind of motion. Indeed, the asymptotic behavior of $\bar{Y}(\gamma(\mathbf{x}_0; t))$ can be summarized in the following way: $\bar{Y}(\gamma(\mathbf{x}_0; t)) \approx a_\gamma t + d_\gamma$, where $a_\gamma = \sigma_\gamma/2$ and $d_\gamma \approx 0$ for irregular, stochastic motion, while $a_\gamma = 0$ and $d_\gamma \approx 2$ for quasi-periodic motion. Deviations from the value $d_\gamma \approx 2$ indicate that γ is close to some particular objects in phase space, being $d_\gamma \lesssim 2$ or $d_\gamma \gtrsim 2$ for stable periodic orbits (or resonant elliptic tori), or unstable periodic orbits (or hyperbolic tori) respectively (see [4] for details). Finally, the quantity $\hat{\sigma}_1 = Y/t$ verifies that

$$\hat{\sigma}_1(\gamma_q(\mathbf{x}_0; t)) \approx \frac{2}{t}, \quad \hat{\sigma}_1(\gamma_i(\mathbf{x}_0; t)) \approx \sigma_i, \quad as \quad t \rightarrow \infty,$$

supporting the fact that in regular domains, $\hat{\sigma}_1$ converges to 0 faster than σ_1 (which goes to zero like $\ln t/t$), while for stochastic domains, both quantities tend to the positive LCN at a rather similar rate.

Let us introduce here a brief comment regarding the computation of the LCN. As it is already well-known, though the definition of the Lyapunov Characteristic Numbers encompasses an integration over an infinite interval of time, their numerical computation involve a rather large but finite time interval and the expected null value corresponding to regular motion is unlikely to be reached. In such a case instead, the 'Finite Time Lyapunov Characteristic Numbers' (FT-LCNs hereafter, following the nomenclature given in [30]) attains a value of order $\ln T/T$, being T the total integration time. Thus, a critical value has to be adopted as 'zero', so that FT-LCNs' values greater or lower than such critical value are regarded as different from or equal to zero respectively.

3. The potential

For the comparison of the MEGNO vs. the FT-LCNs we deal with the potential introduced by [1] which, obtained after the virialization of an N -body self-consistent model composed of one hundred thousand particles, reproduces many features of real elliptical galaxies, such as mass distribution, flattening, triaxiality and rotation (see also, [31]). Nonetheless, it is clear

that a real elliptical galaxy is a much more complex astrophysical system than a purely dynamical one.

This potential seems to provide an adequate scenario for the comparison between the two above mentioned techniques. To this aim, we address the study of the set of randomly selected orbits $\mathbf{O} = \{\mathbf{x}_i(t), i = 1, \dots, 3472, \mathbf{x}_i(0) = \mathbf{x}_{\mathbf{0}_i}\}$, classified by means of the FT-LCNs in [1]. Their initial conditions $\mathbf{x}_{\mathbf{0}_i}$ and their concomitant FT-LCNs values were provided by the authors. We will identify each orbit of the set \mathbf{O} with a label that runs from 1 to 3472.

The equation that reproduces the potential is:

$$V(x, y, z) = -f_0(x, y, z) - f_x(x, y, z) \cdot (x^2 - y^2) - f_z(x, y, z) \cdot (z^2 - y^2), \quad (4)$$

where

$$f_n(x, y, z) = \frac{\alpha_n}{[p_n^{a_n} + \delta_n^{a_n}]^{\frac{ac_n}{a_n}}}, \quad (5)$$

where p_n^2 is the square of the softened radius given by $p_n^2 = x^2 + y^2 + z^2 + \epsilon^2$ when $n = 0$, or $p_n^2 = x^2 + y^2 + z^2 + 2 \cdot \epsilon^2$ for $n = x, z$, and $\alpha_n, \delta_n, a_n, ac_n$ are constants. The adopted value for the softening parameter is $\epsilon \simeq 0.01$ for any n . The functions $f_n(x, y, z)$ were computed through a quadrupolar N -body code for a hundred thousand bodies, which allowed the authors to write them in a general fashion given by Eq. (5). The adopted values for the constants α_n, δ_n, a_n and ac_n are given in Table 1 (further references in [32]).

	α	a	δ	ac
$n = 0$	0.92012657	1.15	0.1340	1.03766579
$n = x$	0.08526504	0.97	0.1283	4.61571581
$n = z$	-0.05871011	1.05	0.1239	4.42030943

Table 1: Adopted values for the coefficients of the functions f_n given by Eq. (5).

The stationary character of the parameters given in Table 1 were tested by performing several fits at different times after virialization, resulting with a precision of 0.1%.

After the system had relaxed, there remained 86.818 bodies resembling an elliptical galaxy (the system obeying a de Vaucouleurs law, as Fig. 2 in [1] shows) with a strong triaxiality and a flattening that increases from the border of the system to its center (see Table I in the same work).

Fig. 1, taken from [32], displays the behavior of the f_n regards to r , being $f_z < 0$, $f_0, f_x > 0$, and $f_x > |f_z|$ for the whole r range, while for $r \gtrsim 0.36$ it is $f_0 > f_x$. Notice that the functions f_n are plotted with the concomitant sign with which they appear in Eq. (4). The obtained triaxial potential has semi-axis X, Y, Z satisfying the condition $X > Y > Z$ and its minimum, which is close to -7 , matches the origin. The potential is less flattened than the mass distribution, as expected (see Table I in [1]).

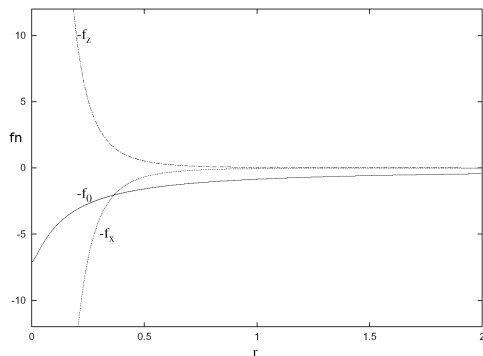


Figure 1: Radial dependency of the functions f_n with their concomitant signs (figure taken from [32]).

4. Comparison of the MEGNO vs. the FT-LCNs

The present section is devoted to performing a numerical and statistical comparative study of the results obtained by recourse to the MEGNO when applied to the set \mathbf{O} of orbits and those provided by [1] through the FT-LCNs.

It is of interest to deem the energy spectrum of the orbits in the set \mathbf{O} displayed in Fig. 2, where we observe that most of the 3472 orbits considered have large energies; indeed, in the main they have energies in the range $-3 \leq E < 0$.

Let us recall the criterion used in [1] to classify the orbits in \mathbf{O} according to their FT-LCNs: those orbits with their largest FT-LCN below some critical value V_c were labeled as regular, otherwise they were classified as chaotic.

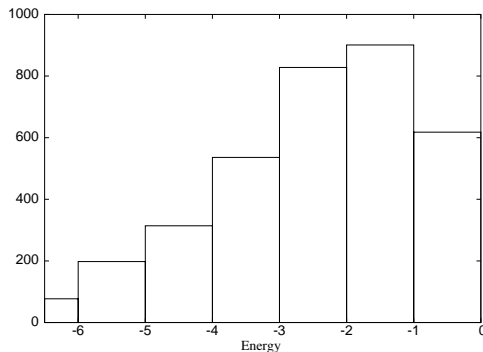


Figure 2: Energy spectrum of the 3472 orbits.

Since the numerical integrations carried out by [1] for the computation of the FT–LCNs encompasses an interval of 10000 u.t. (units of time), the expected value for V_c would be $V_c^t = \ln T/T \approx 0.00092 \text{ (u.t.)}^{-1}$. Notwithstanding, [1] took an empirical value slightly higher, $V_c^e = 0.00155 \text{ (u.t.)}^{-1}$, and this is the one we consider in order to observe their classification into regular and irregular orbits. Along this investigation we adopt for the MEGNO a threshold value of 2.01 for regular orbits. According to [1], the set $\mathbf{O} = \mathbf{O}_c \cup \mathbf{O}_r$, where \mathbf{O}_c and \mathbf{O}_r include 1828 chaotic orbits and 1644 regular ones, respectively.

The computation of the MEGNO, as well as that of the largest FT–LCN, requires the integration of the equations of motion along with their first variationals, the initial conditions for the latter being taken at random in phase space and with unit norm. The integrations were accomplished using a *Runge-Kutta 7/8 th* order integrator (the so-called *DOPRI8* routine –see [33, 34]–), over short: 5000 u.t, intermediate: 10000 u.t. and large integration times: 100000 u.t. The precision in the conservation of the energy was of the order of $\sim 10^{-12}$.

The FT–LCN values corresponding to a total integration time of 10000 u.t. for the sample of orbits classified in [1] were kindly provided by Muzzio, to whom we are grateful.

Herein we present the results corresponding to 5000 u.t. and to 10000 u.t. and even larger motion times, in order to disclose how efficient could this tool be, to provide dynamical information at short times.

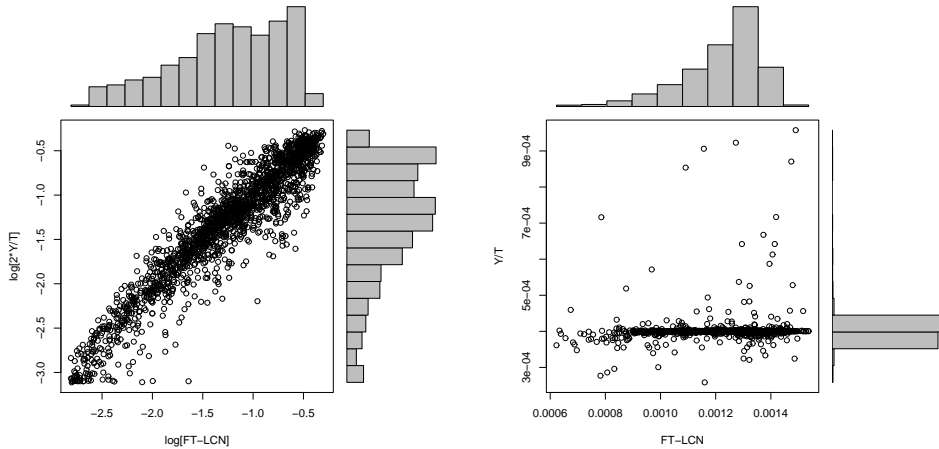


Figure 3: Correlations between $2\bar{Y}/T$ and the largest FT-LCN for chaotic orbits (left panel) and between \bar{Y}/T and the largest FT-LCN for regular orbits (right panel) for $T = 5000$ u.t. The distributions of both the MEGNO and the FT-LCN are also included. These figures were generated using [35].

4.1. Results at short integration times

For the sake of making the comparison clearer, in Figs. 3 and 4 we have changed the scale of \bar{Y} through the division by the total integration time, T , in order to have both the MEGNO and the largest FT-LCN values of the same magnitude. Let us recall that a factor 2 should be added in the case of chaotic orbits since we are dealing with \bar{Y}/T instead of Y/T (see the discussion at the end of Section 2).

At short integration times we observe a few orbits, classified as chaotic by the FT-LCN, falling very close to the regular value of the MEGNO. Indeed, only 0.82% of the subset of chaotic orbits **Oc** attained MEGNO values close to 2.01 (~ -3.095 in Fig. 3, left panel) at $T = 5000$ u.t., while 8.82% of the orbits in the subset **Or** achieved MEGNO values within the range $[2.01, 10)$, indicating either their mild chaotic character or that the total integration time $T = 5000$ u.t. is not large enough for the asymptotic regular value to be reached.

For the orbits in **Oc**, the mean of $\log(\text{FT} - \text{LCN}) \approx -1.267$ and the mean of $\log(2\bar{Y}/T) \approx -1.317$, while the corresponding standard deviations are ≈ 0.617 and ≈ 0.676 , respectively, with a correlation coefficient of $r \approx 0.942$. Thus, both distributions are quite similar.

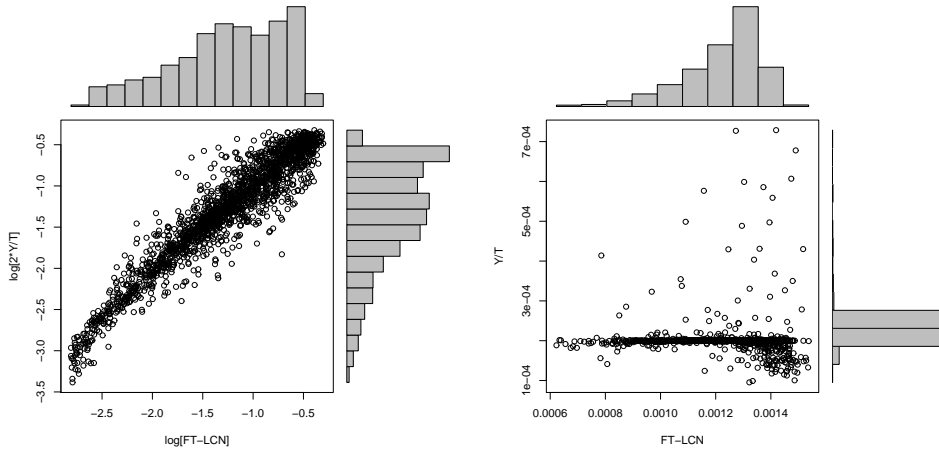


Figure 4: Correlations between $2\bar{Y}/T$ and the largest FT-LCN for chaotic orbits (left panel) and between \bar{Y}/T and the largest FT-LCN for regular orbits (right panel) for $T = 10000$ u.t. The distributions of both the MEGNO and the FT-LCN are also included. These figures were generated using [35].

For the regular sample instead, the concordance is, as expected, less fortunate. The concomitant correlation coefficient r is close to 0.07. The respective mean values are $\text{FT} - \text{LCN} \approx 0.00126$ and $\bar{Y}/T \approx 0.0004$, with rather different standard deviations, namely, 0.00016 for the distribution of the FT-LCNs and 0.000035 for \bar{Y}/T .

Notice must be taken of the fact that we are comparing values of the FT-LCNs and the MEGNO corresponding to different integration times. Altogether, the classification by recourse of the MEGNO provides fairly good results taking account that they are obtained for $T = 5000$ u.t., half the total integration time used by [1] in their computation of the FT-LCN. The comparison of both dynamical indicators at the very same total integration time T is the subject of the forthcoming section.

4.2. Results at intermediate integration times

From Figs. 4 we observe a rather good agreement between the classification given by the MEGNO and that due to the largest FT-LCN at $T = 10000$ u.t. In fact, the chaotic component \mathbf{Oc} appointed by the FT-LCN is re-attained by means of the MEGNO, i.e. all orbits in \mathbf{Oc} have MEGNO values lying on the MEGNO irregularity range (above ~ -3.4 on the vertical axis

in the plot on the left). Since we have rescaled the MEGNO by $1/T$ and $2/T$ for regular and chaotic orbits respectively, those values above, but close to, either 0.0002 or -3.4 after 10000 u.t., suggest that the orbit could be proximate to an hyperbolic object (like unstable periodic orbits, 2D hyperbolic torus). On the other hand, for MEGNO values $\lesssim 0.0002$, the orbit may be close to elliptical objects (stable periodic orbits, 2D resonant elliptical torus), as it is shown in [4].

On the left panel in Fig. 4, corresponding to chaotic orbits, we distinguish a clearly linear correlation between $\log(2\bar{Y}/T)$ and $\log(\text{FT} - \text{LCN})$ for $T = 10000$ u.t. Indeed, the correlation coefficient is $r \approx 0.95$; the mean value of $\log(2\bar{Y}/T) \approx -1.29$, while the concomitant mean value of $\log(\text{FT} - \text{LCN}) \approx -1.27$.

In the same figure on the right, we show the correlation between \bar{Y}/T and the largest FT-LCN for the orbits in **Or**. In this case, the correlation coefficient is $r \approx -0.026$, the mean value of \bar{Y}/T is rather close to 0.0002 with a standard deviation $\lesssim 10^{-5}$, while for the FT-LCNs the mean value is 0.0013 with a standard deviation of order $\approx 10^{-4}$. Let us point out the sharp character of the distribution of the MEGNO values around the predicted one for regular motion, while the FT-LCNs' distribution is rather blunt, as the standard deviation of both distributions indicate. This fact should be deemed as an advantageous feature of the MEGNO over the largest FT-LCN.

Notice must be taken that the empirical value $V_c^e = 0.00155$ adopted by [1] for their classification, is greater than the mean value of the largest FT-LCN for regular orbits, which indeed might be a misleading factor for the task.

Thus, from the regular component **Or**, which encompasses 1644 orbits on the whole, 1513 orbits have values of the \bar{Y}/T in the interval $[0.0001; 0.000201)$, in due accord with their stable, regular character. However, a discrepancy is found for a subset **Or_d**, including the remaining 131 orbits classified as regular by their largest FT-LCN, whose \bar{Y}/T values, however, lie within the range $[0.000201; 0.001)$ revealing their possible irregular character.

An issue to be stressed is the fact that the deemed orbits belong to different energy surfaces and, for each energy, a different characteristic time-scale, $T_c(E)$, can be defined. Moreover, in order to ensure that both the FT-LCN and MEGNO are well computed for a given orbit, the total integration time should verify $T \gg T_c(E)$. Therefrom, on fixing the condition $T \geq 10^3 T_c(E)$ to obtain confident values for both indicators, we conclude that only those orbits with $T_c(E) \leq 10$ would be properly classified for a total integration

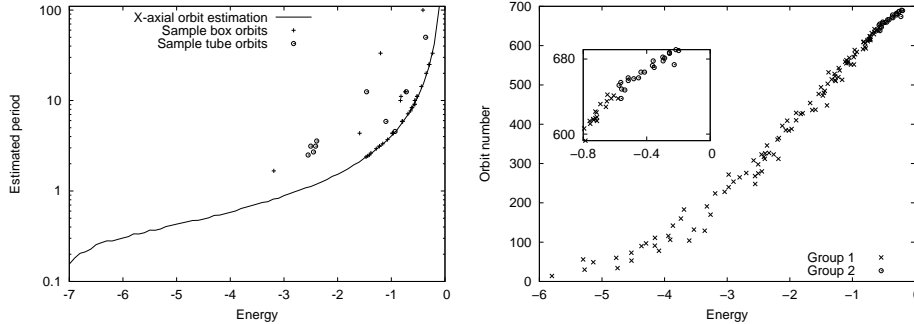


Figure 5: Period of the x-axis orbit as a function of the energy adopted as $T_c(E)$ for box orbits, along with the approximate period of a sample of both box and tube orbits (on the left). Energy of the 131 orbits in \mathbf{Or}_d (on the right).

time of 10000 u.t. The function $T_c(E)$ for box orbits, taken as the period of the stable x-axis periodic orbit, is plotted in Fig. 5 on the left, which shows that $T_c(E) = 10$ corresponds to an energy value $E \sim -0.58$. We have computed the approximate period of some box orbits (crosses in the figure) in order to test the suitability of the adopted time-scale $T_c(E)$ for boxes. The approximate period of some tube orbits are also included in the figure.

The energy values corresponding to the orbits in \mathbf{Or}_d are displayed on the right of Fig. 5, where we have labeled as group 1 those with energies $E < -0.58$, and as group 2 the ones for which $E > -0.58$.

From the 131 orbits in \mathbf{Or}_d 41 are tubes while 90 are boxes, 20 of which have $E > -0.58$, i.e. their dynamical indicators would still be in a transient phase. On the whole, we count 70 box orbits in \mathbf{Or}_d satisfying the condition $T \gg T_c(E)$.

Let us remark that all orbits in \mathbf{Or}_d have $2 \lesssim \bar{Y} \lesssim 7$ at $T = 10000$ u.t. so, even when they could evince some local instability, they behave as stable orbits from a physical point of view. Nonetheless, since our aim is to subject the MEGNO to a rigorous test as a dynamical indicator, we will study this subset of orbits in particular.

Therefore, for the 131 orbits in \mathbf{Or}_d we recalculate the MEGNO but for 100000 u.t. to find that 52 orbits, having MEGNO values smaller than 3.5 at $T = 10000$, approach the regular value 2 at $T = 100000$, while the remaining 79 attain greater values of the MEGNO for the larger integration time. Let us mention that for some orbits the MEGNO value is barely higher than

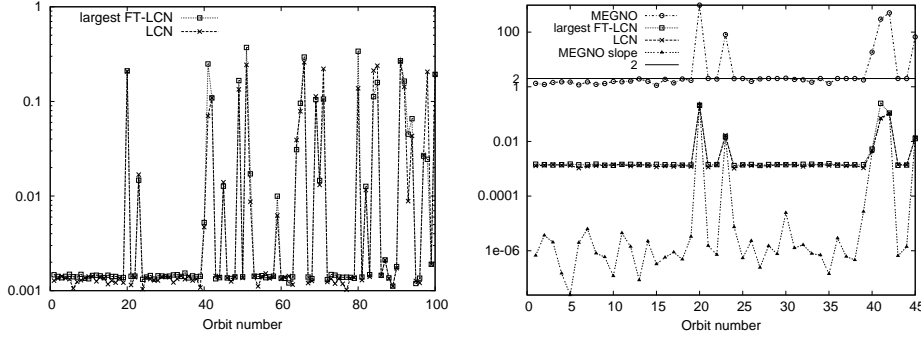


Figure 6: Both the largest FT–LCN provided by [1] and the maximum LCN computed by our using the classical algorithm for a sample of 100 orbits in \mathbf{O} (left). Largest FT–LCN, LCN, slope of the MEGNO and MEGNO for a smaller sample of 45 orbits (right).

the adopted threshold of 2.01 (as it will be shown in section 4.4). In the following subsection, we will identify the first group by \mathbf{Or}_d^s and the second one by \mathbf{Or}_d^u .

In order to determine the actual character of the orbits in \mathbf{Or}_d , we will recourse to a slight variation of the so–called Fast Lyapunov Indicator (FLI) (see [20]). Briefly, the FLI is defined as the supremum of the norm of the tangent vector $\vec{\delta}$. Thus, we will follow the evolution of $\langle \delta(t) \rangle$ where

$$\langle \delta(t) \rangle = \frac{1}{t} \int_0^t \delta(t) dt,$$

with $\delta = \|\vec{\delta}\|$ and $\vec{\delta}$ is the solution of the variational equations, taking as initial value $\delta_0 = 10^{-4}$.

Since motion times larger than 10000 u.t. will be considered in the next section, let us first compute the LCN for a sample of 100 orbits in \mathbf{O} for $T = 10000$ u.t. using the classical algorithm of [36], and compare the obtained values with the largest FT–LCN derived by [1]. The result of the comparison is illustrated in the plot on the left of Fig. 6, which gives account of an actually quite good agreement. This encourages us to compute the LCNs and assimilate them with the largest FT–LCNs for $T = 100000$ u.t. In Fig. 6 on the right we also include the estimation of the largest FT–LCN derived from the slope of the MEGNO and the MEGNO itself for a small sample of orbits (45 on the whole). Let us point out that the slope of the MEGNO yields a better estimation of the largest FT–LCN, particularly for the regular

orbits, for which it lies below 10^{-4} , which is much smaller than both V_c^e and V_c^t for $T = 10000$ u.t. Nonetheless, we will restrict our comparative analysis to the MEGNO, the LCN and the mean FLI ($\langle\delta(t)\rangle$), since deriving the expected theoretical value for the slope of the MEGNO in case of quasi-periodic motion is difficult. In fact, the same occurs with the mean FLI, for which it is not possible to determine an asymptotic value whenever the orbit is confined to a torus. Indeed, for such a quasi-periodic orbit, γ_q , the solution of the variational equation in \mathbb{R}^{2N} can be recast as

$$\delta(\gamma_q(t)) \approx \delta_0 [1 + w_q(t) + t(\lambda_q + u_q(t))],$$

where $\lambda_q > 0$ is the linear rate of divergence around γ_q , and $w_q(t)$ and $u_q(t)$ are oscillating functions of t of bounded amplitude (in general quasi-periodic and with zero average), satisfying $|u_q(t)| \leq b_q < \lambda_q$. The parameter λ_q is a measure of the lack of isochronicity around the orbit since it is related to the maximum eigenvalue of the matrix $\partial\omega/\partial\mathbf{I}$, ω and \mathbf{I} being the frequency and action vectors associated to the torus, respectively (for an isochronous system, such as the harmonic oscillator, $\lambda = b_q = 0$ for all γ).

4.3. Results at large integration times

In the present section we will be concerned with the temporal evolution of the three indicators to be compared, namely the MEGNO, the LCN and the mean FLI (for which no re-normalization was performed and in the case of the exponential growth of $\delta(t)$ the integration was stopped at $\langle\delta(t)\rangle = 10^{20}$), for large motion times. Figs. 7 displays the typical behavior of these indicators for chaotic and regular orbits which are identified by their orbit number. For the illustration the orbits 1491, 442 and 3359 from the set **Oc** and 358, 2105 and 3375 from **Or** – **Or_d** have been selected.

The MEGNO shows a linear growth with time for the orbits in **Oc**, except for orbit 1491 for which exhibits two similar linear trends and a flat behavior between ~ 20000 and ~ 60000 u.t., suggesting that during this time interval the orbit may be close to some elliptic structure. On the other hand, for the orbits in **Or** – **Or_d** the MEGNO asymptotically approaches the predicted value, 2, both at 10000 u.t and 100000 u.t.

In regards to the LCN, a similar behavior is observed for the orbit 1491, while for the regular orbits converges to the theoretical expected value $V_c^t = \ln T/T \sim 1.2 \times 10^{-4}$.

Finally, for those orbits in **Oc** the mean FLI displays an almost exponential dependence with time (in fact it goes as $e^{\sigma t}/t$), while it attains much

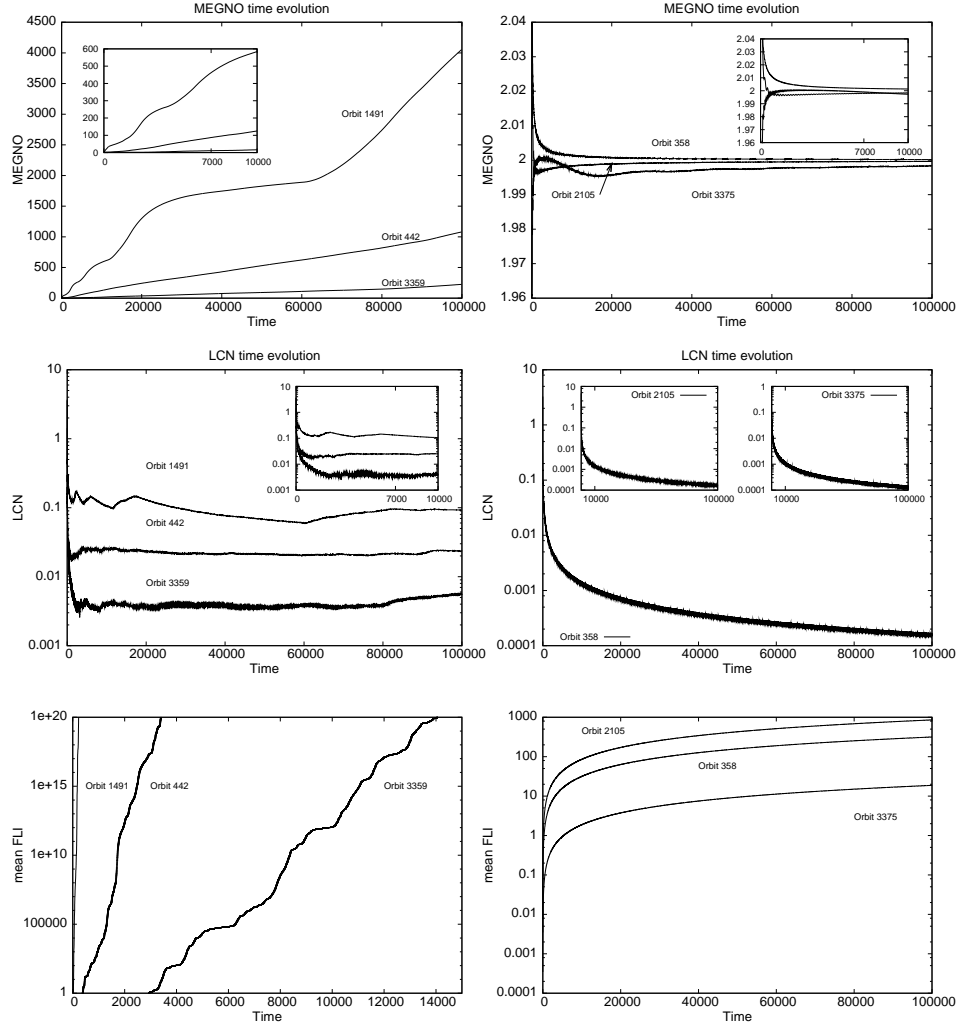


Figure 7: Typical temporal evolution of the MEGNO, the LCN and the mean FLI for chaotic and regular orbits (e.g. in $\mathbf{O}c$ and $\mathbf{O}r - \mathbf{O}r_d$, respectively). For the plots in the last panel, the runs were stopped at $\langle \delta(t) \rangle = 10^{20}$ in order to avoid overflow.

lower values for orbits in $\mathbf{O}r - \mathbf{O}r_d$ for which it depends with time in a linear fashion (note the logarithmic scale in the vertical axis),

Thus, these figures provide information about the expected behavior of the three indicators in the cases of both regular and chaotic motion, which will be of use to determine the character of those orbits in $\mathbf{O}r_d^s$ and $\mathbf{O}r_d^u$.

Let us recall that we will restrict our study to those orbits satisfying the condition $E \leq -0.58$ and for which a good estimation of their period is at hand.

The correlation between the values of MEGNO and the LCN at $T = 100000$ u.t. for orbits in \mathbf{Or}_d^u and \mathbf{Or}_d^s is presented in Fig. 8. For the orbits in \mathbf{Or}_d^u , the mean value of $\log(2\bar{Y}/T) \sim -3.28$ while the mean of $\log(\text{LCN}) \sim -3.13$, and the standard deviations are 0.72 and 0.56 respectively, the correlation coefficient being close to 0.98. Therefore, not only the correlation between both indicators is quite good, but the two first moments of their concomitant distributions are rather similar as well. Let us mention that, though we are computing the MEGNO for very large times, most of the orbits in \mathbf{Or}_d^u attain values in the range $-4.5 \lesssim \log(2\bar{Y}/T) \lesssim -2.5$, revealing that these orbits are mild chaotic, for which the mean Lyapunov time is $T_{Lyap} \sim 1500$ u.t., their characteristic period being smaller than 10 u.t.

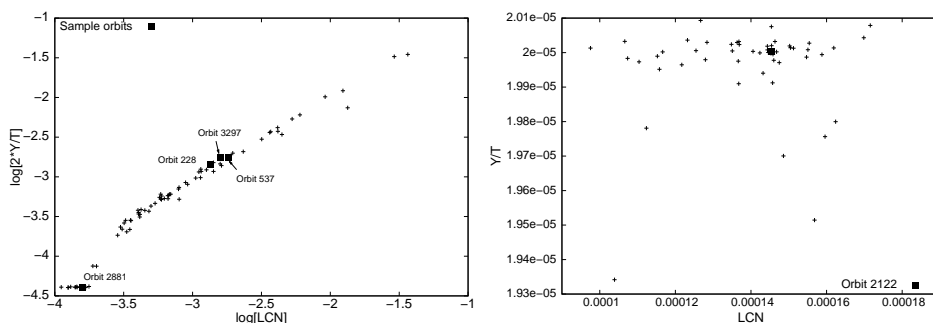


Figure 8: Correlations between $2\bar{Y}/T$ and the LCN for orbits in \mathbf{Or}_d^u (left panel) and between \bar{Y}/T and the LCN for orbits in \mathbf{Or}_d^s (right panel) at $T = 100000$ u.t. The orbits marked with a full dot will be studied separately, following the temporal evolution of the three dynamical indicators.

Meanwhile, for orbits in \mathbf{Or}_d^s , the mean value of $\bar{Y}/T \approx 2 \times 10^{-5}$ with a standard deviation of 1.3×10^{-7} , while the mean LCN is close to 0.00014 with a standard deviation of about 1.7×10^{-5} . The corresponding correlation coefficient is 0.1. Again we point out the sharp distribution of \bar{Y}/T around the expected theoretical value. The standard deviation of both distributions differ in two orders of magnitude (notice should be taken of the different scales onto the vertical and horizontal axis in the right plot of Fig. 8). At $T = 100000$ u.t. we attain values of \bar{Y}/T that provides a fairly good estimation

of the true LCN, $\sigma = 0$, namely, of order 10^{-5} , while the LCN computed by recourse to the standard algorithm barely abuts 10^{-4} .

The explanation is clear; the factor $\ln T \approx 11$ in V_c^t is the responsible for this slower convergence of σ_1 to $\sigma = 0$ as $T \rightarrow \infty$, and \bar{Y}/T tends to σ faster than $\ln T/T$. Indeed, for stable motion and integration times of order of, or larger than $T \approx 22000$ u.t., $|\bar{Y}/T| \lesssim 10^{-4}$, while the LCN computed using the standard algorithm yields $|\sigma_1| \lesssim 10^{-3}$. In fact,

$$\frac{\bar{Y}/T}{\sigma_1} \approx \frac{2}{\ln T} \rightarrow 0, \quad T \rightarrow \infty.$$

Finally, the full dots in Fig. 8 correspond to five orbits selected as samples of \mathbf{Or}_d^s and \mathbf{Or}_d^u , for which the study of the temporal evolution of the three indicators for $T = 100000$ u.t. will serve to determine their dynamical behavior. This issue will be undertaken in the forthcoming subsection.

4.4. On the \mathbf{Or}_d orbits

Let us be concerned with the detailed study of some sample orbits of \mathbf{Or}_d . First we will aim our attention at orbits 2881, 537 and 3297 belonging to \mathbf{Or}_d^u , and 2122 $\in \mathbf{Or}_d^s$. The study of orbit 228 $\in \mathbf{Or}_d^u$ will be addressed separately.

Fig. 9 displays the temporal evolution of the three indicators for both integration times, namely, $T = 10000$ and $T = 100000$ u.t., corresponding to the selected sample orbits. It can clearly be observed that for $T = 10000$ u.t. almost all orbits exhibit a stable behavior. Yet, both the MEGNO and the mean FLI evince an incipient increase for orbits 537 and 3297, which is missed by the LCN. As time increases, both orbits clearly separate from the rest, all the three indicators giving account of this fact.

Meanwhile, orbits 2122 $\in \mathbf{Or}_d^s$ and 2881 $\in \mathbf{Or}_d^u$ seem to evolve in a similar fashion. Though, the final MEGNO value for orbit 2122 is sharply 2, while for 2881 is slightly above the regular value, which might indicate a rather mild unstable character of this orbit (see discussion below).

In Fig 10 we present the time evolution of both the MEGNO and the LCN, on the left and right panel respectively, for orbit 228 $\in \mathbf{Or}_d^u$. It is interesting to note the particular behavior of the indicators for this orbit. From the plot on the left, the trajectory looks like a stable quasi-periodic orbit up to $t \lesssim 2000$ u.t., then the MEGNO grows linearly for a rather short time interval to reach a nearly constant value, around 2.6, and attains the

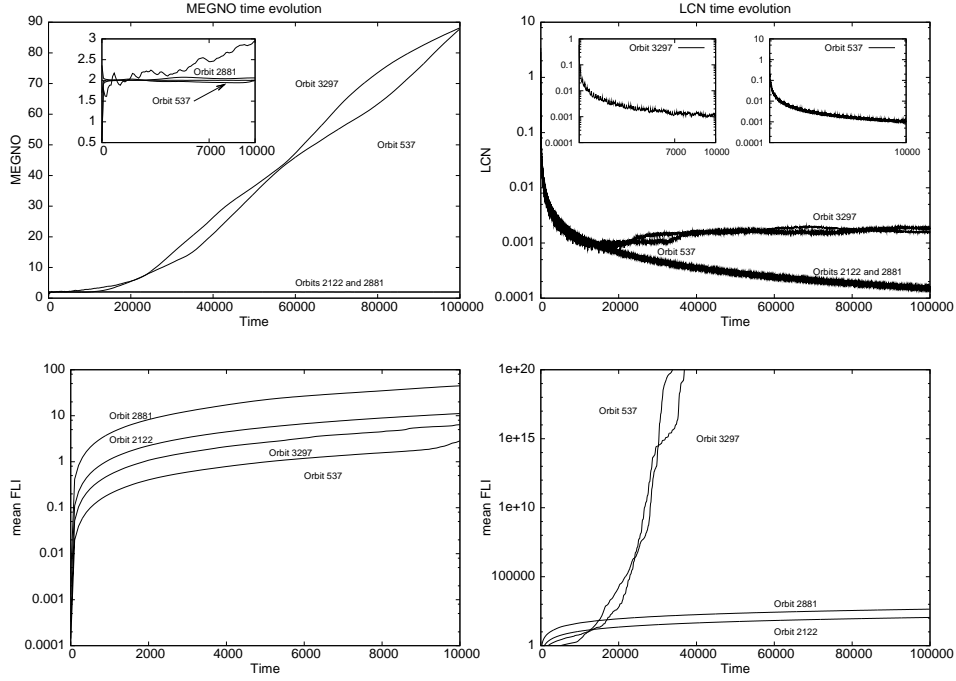


Figure 9: Time evolution of the MEGNO, the LCN and the mean FLI for the sample orbits in \mathbf{Or}_d^u and \mathbf{Or}_d^s for $T = 10000$ and $T = 100000$ u.t. In the bottom panels we separately plot the evolution of the mean FLI for different times.

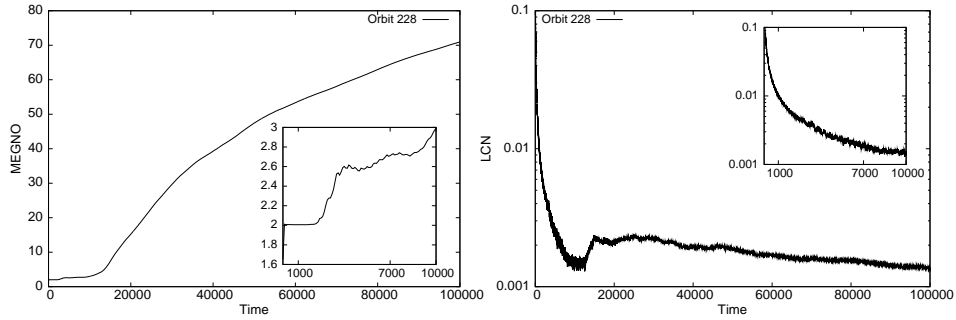


Figure 10: Time evolution of the MEGNO and the LCN for orbit 228 $\in \mathbf{Or}_d^u$ for $T = 10000$ and $T = 100000$ u.t.

value 3 at $T = 10000$ u.t. Note that the LCN at this time is very close to

the theoretical expected one, around 0.001. When the integration time is increased the MEGNO grows up to higher values but not in a linear fashion, while the LCN seems to decrease, though it approaches a larger value than the one corresponding to regular motion at $T = 100000$ u.t. The peculiar behavior of the indicators for this orbit encourages a more detailed study of its neighbourhood in phase space in order to grasp its actual dynamical nature. This chore will be performed by analysing its immediate neighbourhood in phase space, which might provide us with valuable dynamical information.

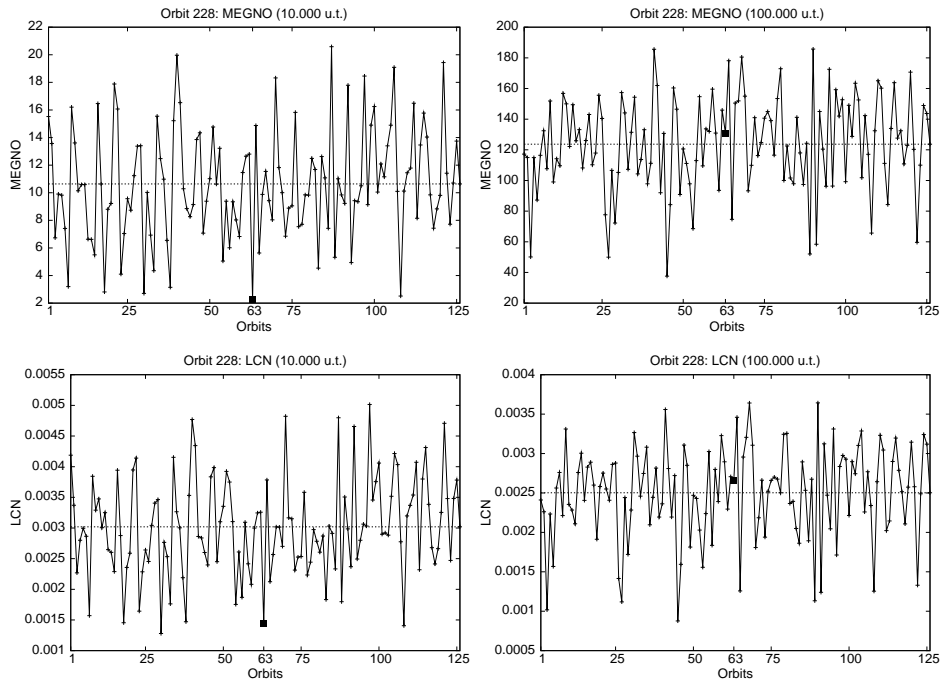


Figure 11: MEGNO and the LCN for $T = 10000$ and $T = 100000$ u.t. for 125 orbits in a domain of size 10^{-7} around orbit 228. The last value corresponds to the mean value of the indicators, also represented by the horizontal line.

Thus, in Fig. 11 we present the MEGNO and LCN values at $T = 10000$ and $T = 100000$ u.t. for a set of 125 orbits taken at random in a neighborhood of size 10^{-7} centered at orbit 228, whose concomitant values are depicted by full dots in each plot. Note that at $T = 10000$ the MEGNO for this orbit is very close to 2, while the mean value of the indicator for this set of orbits is about 11. On increasing the integration time, it becomes quite clear that

orbit 228 is in fact chaotic. A similar behavior is observed in regards to the LCN values. Altogether, the figure suggests that this orbit might lie in a complex dynamical region of phase space.

In order to confirm this conjecture, in Fig. 12 we show a MEGNO contour plot in the momenta space for the exact energy value of orbit 228, constructed by taking as initial conditions the position of the orbit in configuration space and (p_x, p_z) varying over a grid of about 10^6 points. The MEGNO values correspond to $T = 1000$ u.t., which turns out to be a proper final time of integration since $T \approx 10^3 T_c(E)$, and the characteristic time scale is about 1 for an energy $E \sim -3$ (as follows from Fig. 5).

The dark regions corresponds to strong chaotic motion for which $\overline{Y} > 20$, while the white ones with $\overline{Y} < 2.01$ reveal stable motion. Light gray zones refer to slight unstable (or even regular) motion, $2.01 < \overline{Y} < 3$ and the dark gray regions represent mild chaotic motion $3 < \overline{Y} < 20$. This plot reveals the complex resonance structure of phase space when projected onto the plane (p_x, p_z) at this energy level, and we can clearly see that orbit 228 lies inside a resonance crossing. This should explain the pathological behavior of this orbit and its surroundings.

Just to end this section, let us investigate the surroundings of orbit 2881 which belongs to \mathbf{Or}_d^u . For that sake, let us consider 125 orbits selected at random within a neighbourhood of size 10^{-7} and compute both their MEGNO and LCN at $T = 10000$ and $T = 100000$ u.t. The results are displayed in Fig. 13 to show that the MEGNO values lie in the range $1.95 \lesssim \overline{Y} \lesssim 3$ for $T = 10000$, and on considering larger motion times the MEGNO interval gets even narrower, e.g. $(1.98, 2.27)$ for $T = 100000$ u.t. Therefore, though the MEGNO for orbit 2881 is slightly higher than the threshold 2.01, namely, 2.014 at $T = 100000$, the orbit should be considered at all means stable. An analogous result provides the LCN, that has been computed for the very same orbits, including orbit 2881 for which, at $T = 10000$ u.t., attains a value less than the empirical critical value V_c^e adopted by [1], but higher than V_c^t .

Finally, it would be interesting to consider the MEGNO values of all orbits in \mathbf{Or}_d^u for $T = 100000$ u.t., which are displayed in Fig. 14. It can there be noticed that at least 14 orbits in \mathbf{Or}_d^u should actually be included in the set \mathbf{Or}_d^s (e.g. 2881), since their MEGNO values are rather too close to the regular value 2. In fact, it might be inaccurate for them to consider the factor 2 in $2\overline{Y}/T$, necessary in the case of chaotic orbits, on looking for correlations with the LCN, since these orbits do not increase linearly with time.

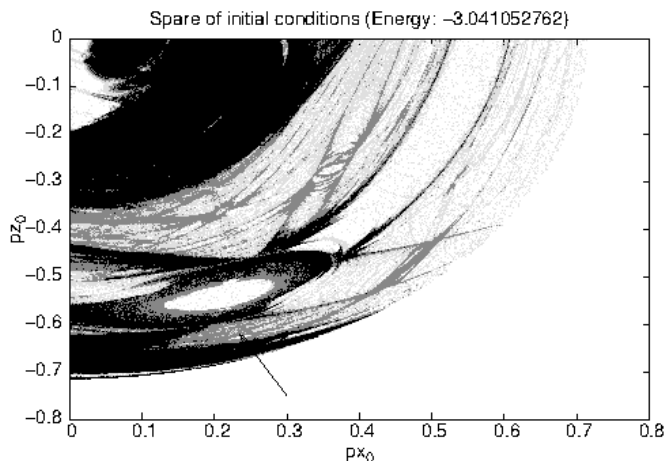


Figure 12: MEGNO contour plot for a grid of 10^6 initial conditions in momenta space (p_x, p_z) for $E = -3.041052762$ and initial position of orbit 228. The total motion time is $T = 1000 \approx 10^3 T_c(E)$. Black ($\bar{Y} > 20$) corresponds to strong chaotic zones, white indicates regular regions $\bar{Y} < 2.01$, light gray to slight unstable or even some regular orbits ($2.01 < \bar{Y} < 3$) and dark gray indicates mild chaotic domains ($3 < \bar{Y} < 20$). The arrow indicates the location of orbit 228.

On transferring these 14 orbits from \mathbf{Or}_d^u to \mathbf{Or}_d^s , the recomputation of the concomitant resulting distributions deliver for \mathbf{Or}_d^u , the mean value of $\log(2\bar{Y}/T) \approx -3.04$ with a standard deviation of about 0.55 and the mean of $\log(\text{LCN}) \approx -2.98$ with a standard deviation close to 0.51, the correlation coefficient being $r \approx 0.99$. For \mathbf{Or}_d^s there results a mean value of $\bar{Y}/T \approx 2 \times 10^{-5}$ with a standard deviation of $\approx 2.3 \times 10^{-7}$, while the mean LCN is $\approx 1.4 \times 10^{-4}$ with a standard deviation of $\approx 1.8 \times 10^{-5}$. The correlation coefficient in this case barely amounts 0.22, indicating again no correlation between both indicators for regular orbits. Let us notice that no significant changes arise as a consequence of the transposition performed.

5. Discussion

We have shown that the MEGNO is a suitable fast indicator to separate regular from chaotic motion. Further, it is particularly useful to investigate the nature of orbits that have a small but positive Lyapunov number.

Besides we have shown a rather good correlation between the MEGNO and the FT-LCN values for short, moderate and large integration times for

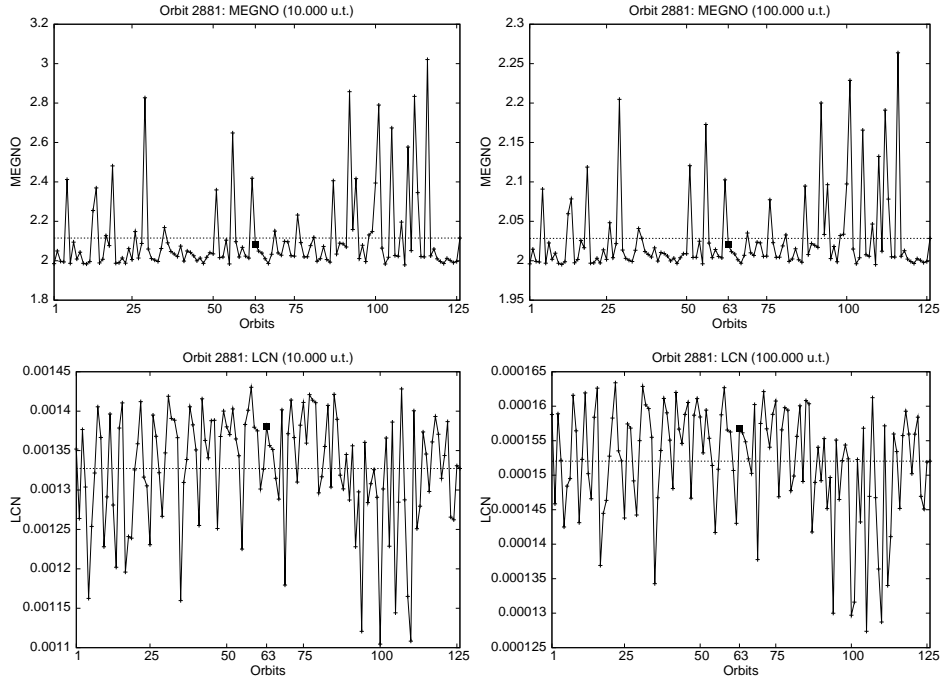


Figure 13: MEGNO and the LCN for $T = 10000$ and $T = 100000$ u.t. for 125 orbits in a domain of size 10^{-7} around orbit 2881. The last value corresponds to the mean value of the indicators, also represented by the horizontal line.

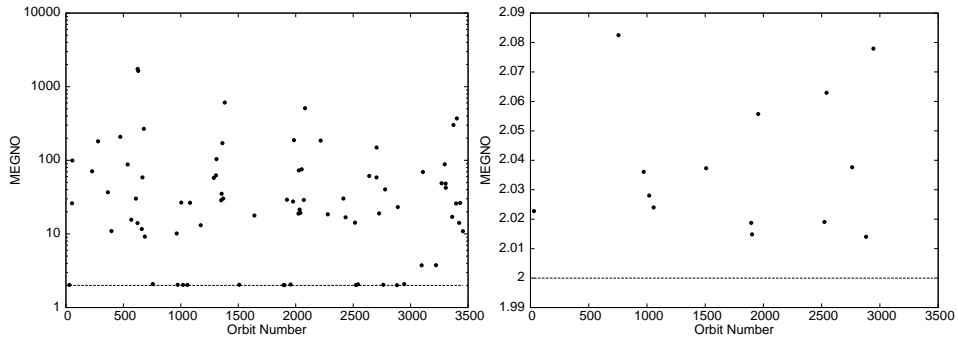


Figure 14: MEGNO values for the 79 orbits in \mathbf{Or}_d^u for $T = 100000$ (left). Zoom around $\bar{Y} = 2$ (right).

chaotic orbits, while the MEGNO provides better results for regular motion.

In fact, it has the advantage that \bar{Y}/T converges to the null value of σ faster than the classical algorithm to compute the LCN. Another recourse to derive low values for the LCN in the case of quasi-periodic motion consists in computing the slope of the MEGNO.

The FLI looks also as a reliable fast indicator, but it does not provide any reference value for regular motion, so it may be useful to explore phase space rather than to investigate the nature of a given orbit, unless we follow the time evolution of this indicator.

Finally, though $\bar{Y} \rightarrow 2$ when $t \rightarrow \infty$ and for quasi-periodic motion its convergence would be very fast, one should keep in mind that a single orbit is in general dominated by the dynamics of its surroundings. Therefore if it is necessary to determine the strict character of that orbit, the threshold value of a regular orbit would be $\bar{Y} \sim d$ where d could be taken as the mean value of \bar{Y} over a rather small domain around the orbit.

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